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UNIVERSITÀ
DEGLI STUDI
DI TORINO

On the origin of the Giant Gamma-ray Halo of M31

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Overview

- summary of gamma-ray observations of M31
- gaseous halos of Milky Way and M31

- **giant gamma-ray halo of M31**
  - energetics
  - possible cosmic-ray origin

- M31 and MW: gamma-rays and neutrinos

- summary
- perspectives & issues

Gamma-ray observations - M31

PAST

- Fermi-LAT - integrated gamma-ray luminosity

  $$L_\gamma(> 100\text{MeV}) \approx 6.6 \times 10^{41}\ \text{erg/s}$$

  Abdo et al. 2010

- concentrated in the **inner ~ 5 kpc region**
- does **not correlate with the gaseous disk**

  Ackermann et al. 2017

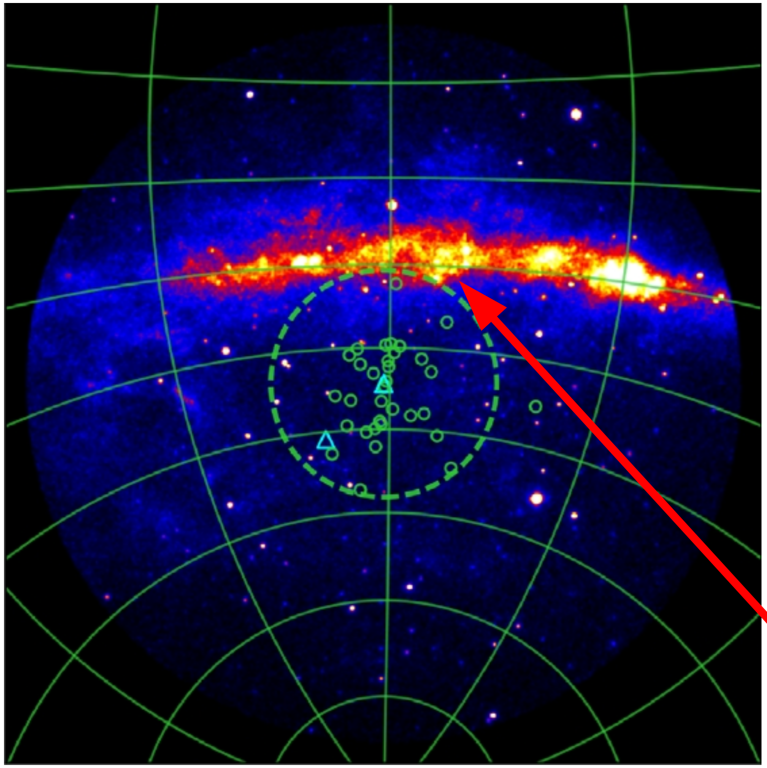
- evidence for ***Fermi Bubbles***-like structure

  Pshirkov et al. 2016

Gamma-ray observations - M31

RECENT ANALYSIS OF *Fermi*-LAT DATA

Karwin et al. 2019



Contamination with MW disk

Photons 1-100 GeV

- 28 deg x 28 deg region
- **200 kpc** projected radius
- excess up to ~ 120-200 kpc

Spherical template

- IG $r < 5.5$ kpc
- SH $5.5 < r < 120$ kpc
- OH $120 < r < 200$ kpc

+ **North** and **South** arcs

Gamma-ray observations - M31

Karwin et al. 2019

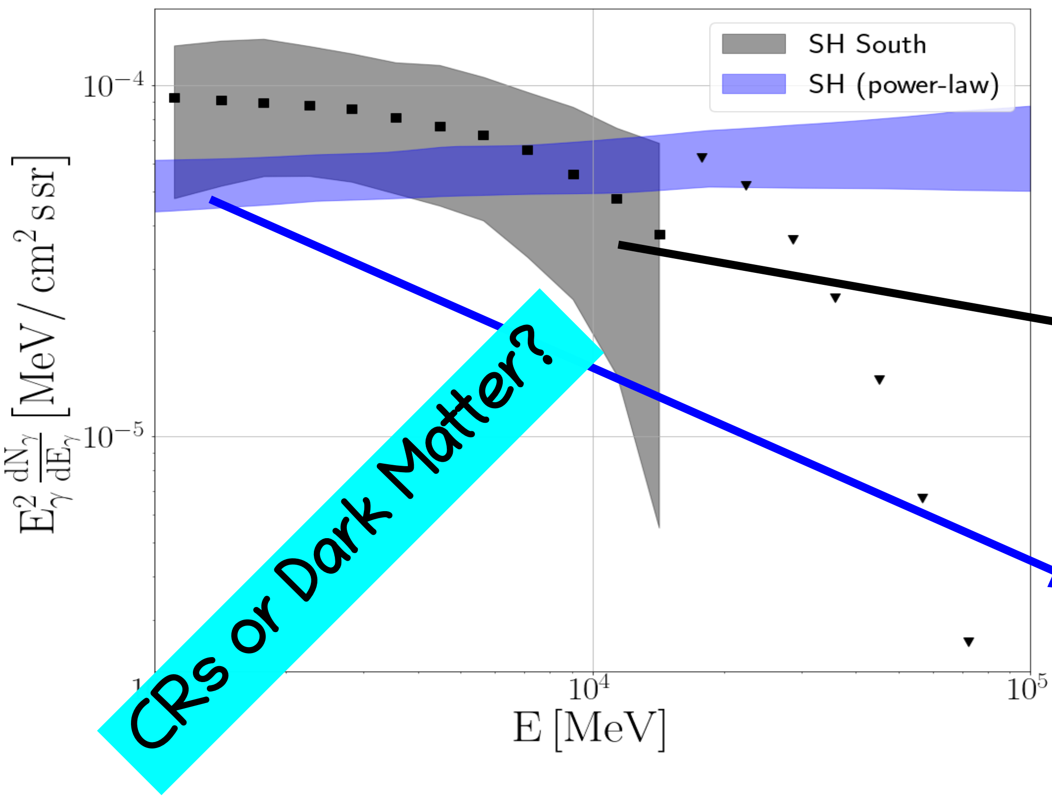
Photons 1-100 GeV

SH likely connected with the M31

for OH not clear

$$L_\gamma^{SH} \sim 1.7 - 1.9 \times 10^{39}\,\mathrm{erg/s}$$

The spectrum is unclear...



Power-law + cut off (best fit)

$$\propto E_{\mathrm{GeV}}^{-1.9} \exp(-E_{\mathrm{GeV}}/11.6)$$

Power-law (other analysis)

$$\propto E_{\mathrm{GeV}}^{-2.0}$$

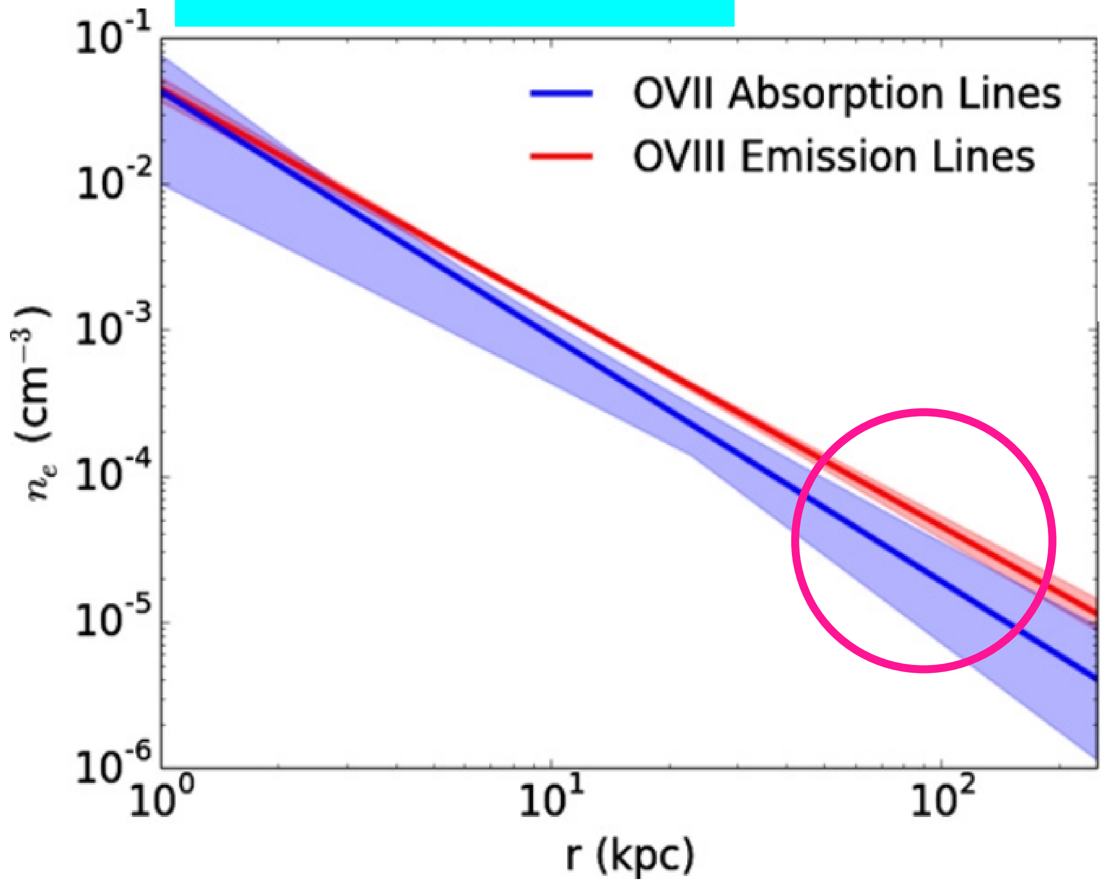
Gaseous Halos MW and M31

- detected extended halo (~ 200 kpc) around the MW
  - OVII, OVII X-ray abs/em lines — Miller & Bregman 2013, 2015; Fang et al. 2012; Gupta et al. 2012
  - ram pressure stripping dwarf — Gatto et al. 2013
  - cosmological simulations — Nuza et al. 2014
- $n \sim 10^{-4}$–10^{-3} cm^{-3}, $T \sim 10^6$, $M \approx 10^{10} - 10^{11}\, M_\odot$
- hot bridge between MW and M31 — Qu et al. 2020
- missing barions, gamma-rays, neutrinos

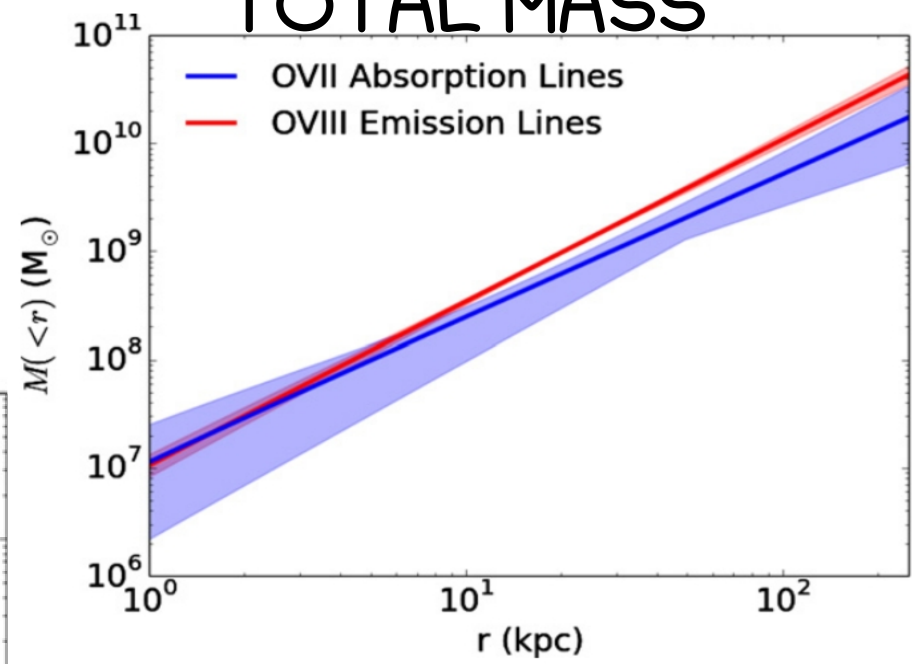
Gaseous Halos MW and M31

Miller & Bregman 2015

GAS DENSITY



TOTAL MASS



SH emission: energetics

$$L_\gamma^{SH} \sim 1.7 - 1.9 \times 10^{39}\, \mathrm{erg/s}$$

Recchia et al. 2019

$$E_e \approx 0.6 - 6\, \mathrm{TeV}$$

Leptonic scenario

$$\tau_{CMB} \sim 1.3 \times 10^6 E_{\mathrm{TeV}}^{-1}\, \mathrm{yr}$$

$$L_e = L_\gamma^{SH}$$

$$E_p \approx 10 - 1000\, \mathrm{GeV}$$

Hadronic scenario

$$\tau_{pp} \sim 7 \times 10^{10} n_{H,-3}^{-1}\, \mathrm{yr}$$

$$\tau_{res} \approx 10^9 - 10^{10}\, \mathrm{yr}$$

$$n_H \sim 10^{-4} - 10^{-3}\, \mathrm{cm}^{-3}$$

$$L_{CR}^{MW} \sim 10^{41}\, \mathrm{erg/s}$$

$$\omega_{CR}^{MW,disk}(1\, \mathrm{GeV}) \sim 1\, \frac{\mathrm{eV}}{\mathrm{cm}^3}$$

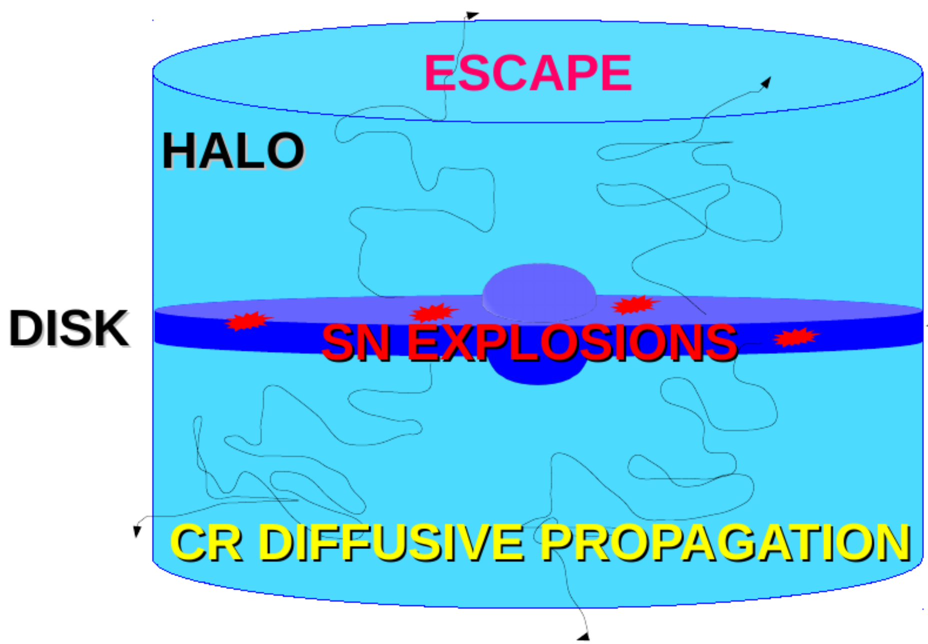
$$\propto E^{-0.7}$$

$$L_p \approx 1.8 \times 10^{41}\, \tau_{res,9}^{-1} n_{H,-3}^{-1}\, \mathrm{erg/s}$$

$$\omega_{CR,p} \approx 0.02\, n_{H,-3}^{-1}\, \mathrm{eV/cm}^3$$

Typical CR transport scenarios

- source(s) in the GC or disk
- "direct" diffusion (+ advection ...) toward halo



MAY IT WORK?

Obvious for electrons (losses)

Typical CR transport scenarios

protons

$$\epsilon(r) \propto n_{CR}(r) n_H(r)$$

emissivity — CR density — gas density

Optimistic scenario, moderate drop of CR density

- Constant D and continuous injection $n_{CR}(r) \propto 1/r$
- Gas density drops with r $n_{disk} \sim 1\,\mathrm{cm}^{-3}$, $n_{SH} \sim 10^{-4} - 10^{-3}\,\mathrm{cm}^{-3}$

$$\frac{\epsilon(r=1)}{\epsilon(r=100)} \approx 10^5 \frac{n_{disk}}{n_{SH,-3}}$$

$$\frac{\epsilon(IG)}{\epsilon(SH)} \approx 10^3$$

With D(r), a galactic wind ..., even worse

... or, if we fix the SH emissitivity, the M31 disk should be very bright in gamma rays

Way out 1: in situ acceleration

Acceleration of e- and p at a **shock** located in the SH, R_{SH}~**100 kpc**

Recchia et al. 2019

$$L_p \approx 1.8 \times 10^{41}\, \tau_{res,9}^{-1} n_{H,-3}^{-1}\, \mathrm{erg/s}$$

$$L_\gamma^{SH} \sim 1.7 - 1.9 \times 10^{39}\, \mathrm{erg/s}$$

$$L_{SNR}^{MW} \sim 10^{42}\, \mathrm{erg/s}$$

accretion shock (existence in M31 or MW?)

$$\nu_{ff} \sim 0.3 \times 10^3 M_{12}^{1/2} R_{SH,2}^{-1/2}\, \mathrm{km/s}$$

free fall velocity

$$L_s \approx (4\pi R_{SH,2}^2) \frac{\rho_0 \nu_{ff}^2}{2}$$

$$\approx 3.4 \times 10^{42} M_{12}^{3/2} R_{SH,2}^{1/2} n_{0,-4}\, \mathrm{erg/s}$$

shock luminosity

GALAXY MASS

acceleration efficiency < few %

Way out 1: in situ acceleration

Acceleration of e- and p at a **shock** located in the SH, R_{SH}~**100 kpc**

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$$L_\gamma^{SH} \sim 1.7 - 1.9 \times 10^{39}\, \mathrm{erg/s}$$

$$L_{SNR}^{MW} \sim 10^{42}\, \mathrm{erg/s}$$

termination shock powered by the GC activity

$$u_s \approx 0.2 \times 10^3 L_{GC,43}^{1/5} \tau_{GC,9}^{-2/5} n_{0,-4}^{-1/5}\, \mathrm{km/s}$$

shock velocity

$$L_{Edd}\ 1.3 \times 10^{46} \left(\frac{M_{BH}}{10^8\, M_\odot} \right) \mathrm{erg/s}$$

Eddington luminosity

BLACK HOLE MASS

Way out 1: in situ acceleration

CR acceleration $\tau_{acc} \sim aD(E)/u_s^2$ $a \sim 10$

electrons	protons
$E_e \approx 0.6 - 6\,\mathrm{TeV}$	$E_p \approx 10 - 1000\,\mathrm{GeV}$
$\tau_{acc} = \tau_{loss}$	$\tau_{acc} = \tau_{res}$
$E_{max,e} \approx 24\, u_{s,3}^3 n_{0,-4}^{1/4}\,\mathrm{TeV}$	$E_{max,p} \approx 460\, u_{s,3}^3 n_{0,-4}^{1/2}\,\mathrm{PeV}$

Efficiencies < few %

No problem with Emax and efficiency,for both e- and p

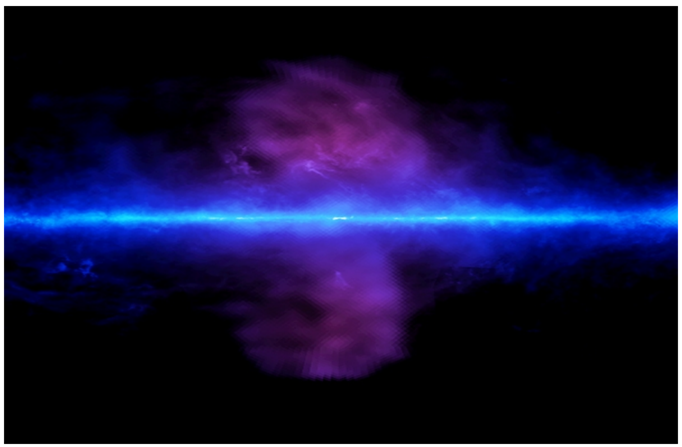
Spectral slope depends on the shock Mach number...,
$T \sim 10^6$, $c_s \sim 100$ km/s, weak shocks?

Way out 2: buoyant bubbles

Recchia et al. 2019

Fermi -Bubbles
similar structure in M31

Pshirkov et al. 2016



$r \gtrsim 10\,\mathrm{kpc}$

star formation/ GC activity

timescale few-few tens Myr

$10^{41} - 10^{43}\ \mathrm{erg/s}$

$W_B \approx 10^{55} - \mathrm{few}\, 10^{57}\ \mathrm{erg}$

Buoyant Bubbles

- Often present in central regions of galaxy clusters (**in galaxies?**)
- Radius of few kpc
- Rise velocity ~ sound speed (~ 100 km/s, HIM)
- Lifetime $\tau_b \approx 10^8\ \mathrm{yr}$
- Stabilizing action of a B-field $\tau_b \gtrsim 10^9\ \mathrm{yr}$

$W_B \approx 10^{57},\ \mathrm{up\ to}\ 10^{59}\ \mathrm{erg}$

Fermi Bubbles base of a larger structure ?

Way out 2: buoyant bubbles

Buoyant Bubbles

- **Periodic activity of the M31 GC emits bubbles** $\nu_b \approx 1/10^8\,\mathrm{yr}^{-1}$
- Inject CRs with average luminosity $L_p =\sim 3.2 \times 10^{41} \eta E_{B,57} \nu_{B,-8}\,\mathrm{erg/s}$
- Efficiency (compare with required luminosity) takes into account adiabatic energy losses $\eta \approx 0.56 \tau_{res,9}^{-1} n_{H,-3}^{-1} E_{B,57}^{-1} \nu_{B,-8}^{-1}$

- CRs are **transported within a bubble into the SH halo**
- With typical **rise velocity $c_s \sim$ 100 km/s** before disruption
- **After disruption** CR diffuse spherically ~ 100 kpc distance

$$\tau_{rise} \approx 10^9\,\mathrm{yr}\frac{R_{\mathrm{SH},2}}{c_{s,2}}$$

$$\tau_{diff} = R_{SH}^2/6\,D \approx 10^9\,\mathrm{yr}\frac{R_{\mathrm{SH},2}^2}{D_{30}}$$

Way out 2: buoyant bubbles

Buoyant Bubbles

- CRs are **transported within a bubble into the SH halo**
- With typical **rise velocity $c_s \sim$ 100 km/s** before disruption
- **After disruption** CR diffuse spherically ~ 100 kpc distance

$$\nu_B > 1/\tau_{diff}$$ → Contribution of several bubbles overlaps - stationarity

$$\nu_B < 1/\tau_{diff}$$ → Intermittency

$$E_c \approx 100\,\mathrm{GeV}$$

Cut-off, observed ?

Constrain on D

$$D(E) < D(E_c) = 5 \times 10^{30} R^2_{SH,2} \nu_{B,-8}\,\mathrm{cm}^2/\mathrm{s}$$

Way out 2: buoyant bubbles

Buoyant Bubbles - CARTOON

- if D flat in energy
- or CRs confined for the lifetime of M31

→ No Ec

Few kpc radius

$R \sim 100\,\mathrm{kpc}$

GC

SH

Bubble disrupt diffusion

GC

... summary so far

The extended gamma-ray emission of M31 can accounted for by

GC activity
(buoyant bubble, termination shock) → Scale with BH mass

Accretion shock
→ Scale with galaxy mass

- in both scenarios the **CR population in the disk and in the halo are decoupled**
- at high energies (> TeV) CR **density in halo can be > in disk** (hard vs steep spectrum)
- this is also important for neutrinos from the MW halo
- energetics, morphology and spectrum can be accounted for

Diffuse Icecube neutrinos from the MW halo

Seminal work by Taylor et al. 2014

$$\frac{N_\nu^{\rm halo}}{N_\nu^{\rm disk}} \approx 0.5 \frac{n_{-3}^h}{n_0^d} \frac{R_2^h}{R_1^d}$$

- relative number of neutrinos from disk and halo with uniform CR population
- **hard ~ E^{-2}** CR spectrum in the halo
- large confinement times

$$L_\nu = K_\nu \left(1 - e^{-\frac{\tau_{\rm res}}{\tau_{\rm pp}}}\right) L_p$$

neutrino VS proton luminosity

$$\tau_{pp} \sim 7 \times 10^{10} n_{H,-3}^{-1}\,{\rm yr}$$

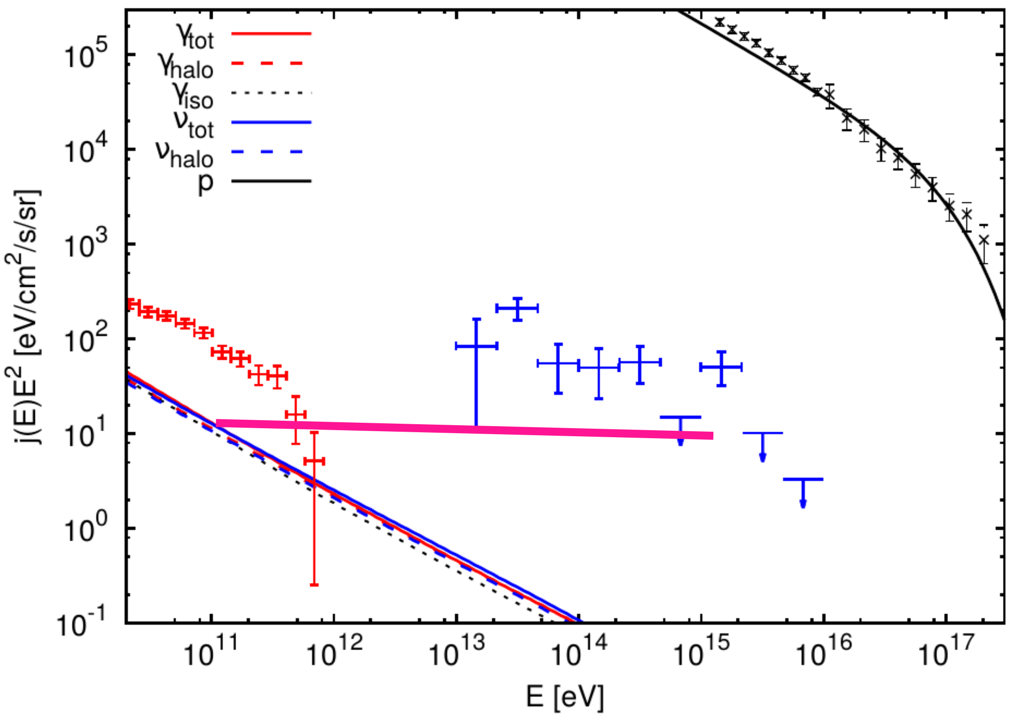
$$\tau_{res} = 3 \times 10^9\,{\rm yr} \frac{R_{\rm SH,2}^2}{D_{30}} = 3 \times 10^9\,{\rm yr} \frac{R_{\rm SH,2}}{u_{\rm adv,30}}$$

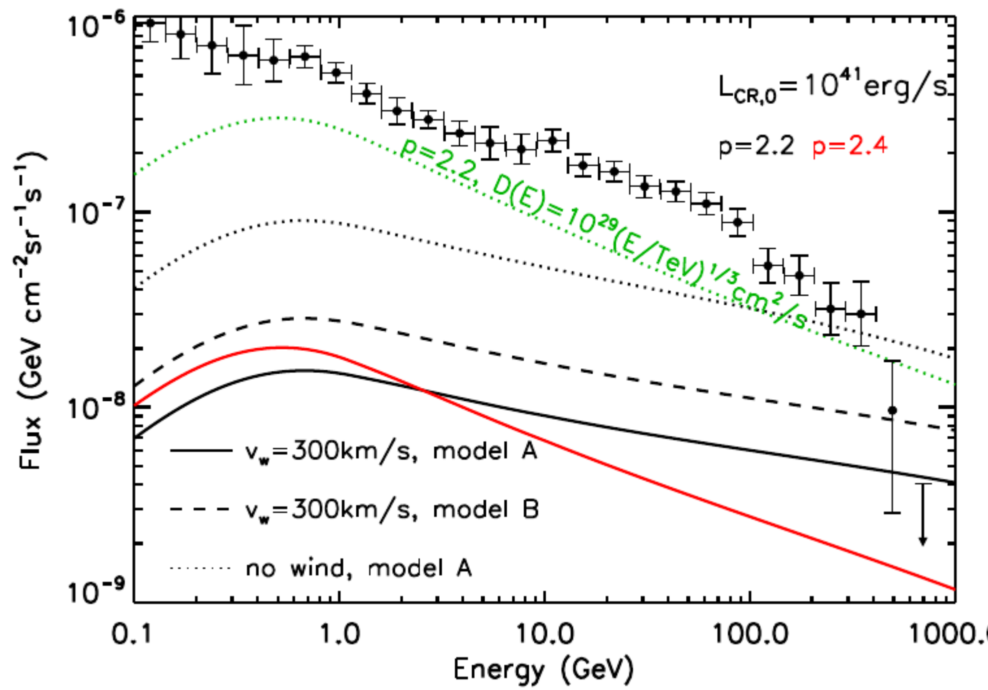
R~ 100 kpc and $t_{pp} < t_{res}$ gives **$L_\nu \sim L_p \sim 10^{39}$ erg/s**
neutrino spectrum reflects source spectrum

Diffuse Icecube neutrinos from the MW halo

Kalashev & Troitsky 2016

Liu et al. 2019 (gamma-rays)





Injection in GC ~ $E^{-2.4}$
D(E) ~ 10^{29} cm^2/s $E^{1/3}$

Need hard spectrum in halo
Need to "decouple" disk-halo

Self-confinement in MW halo

Blasi & Amato 2019

- current of CRs trying to free-escape from the disk in regions with low-B
- excitation of Bell instability
- suppression of diffusion coefficient
- displacement of plama at v_A
- 10^8-10^9 yr in ~ 10 kpc halo

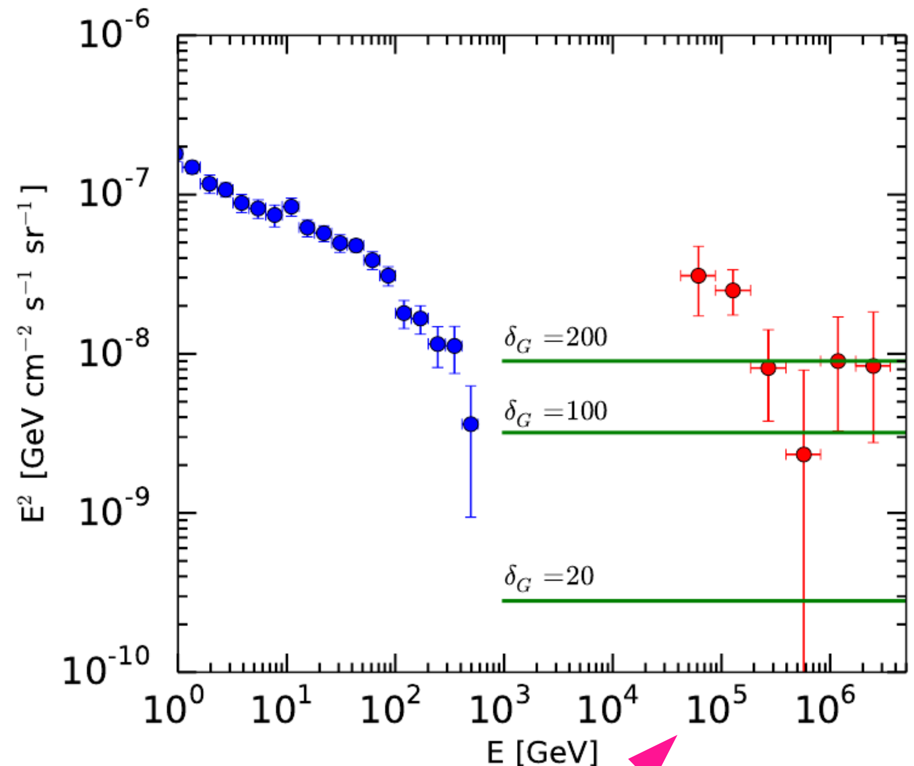
$$\frac{E^2 \phi_{CR}}{c} > \frac{B_0^2}{4\pi}$$

$$\phi_{CR}(E) = \frac{L_{CR}}{2\pi R_d^2} E^{-2}$$

$$B_{\rm sat} \approx 2.2 \times 10^{-8} L_{41}^{1/2} R_{10}^{-1} {\rm G}$$

$$D(E) \approx 1.5 \times 10^{24} E_{\rm GeV} L_{41}^{1/2} R_{10} {\rm cm}^2/{\rm s}$$

$$v_A \approx 5 \times L_{41}^{1/2} R_{10}^{-1} n_{-4}^{-1/2} {\rm km/s}$$



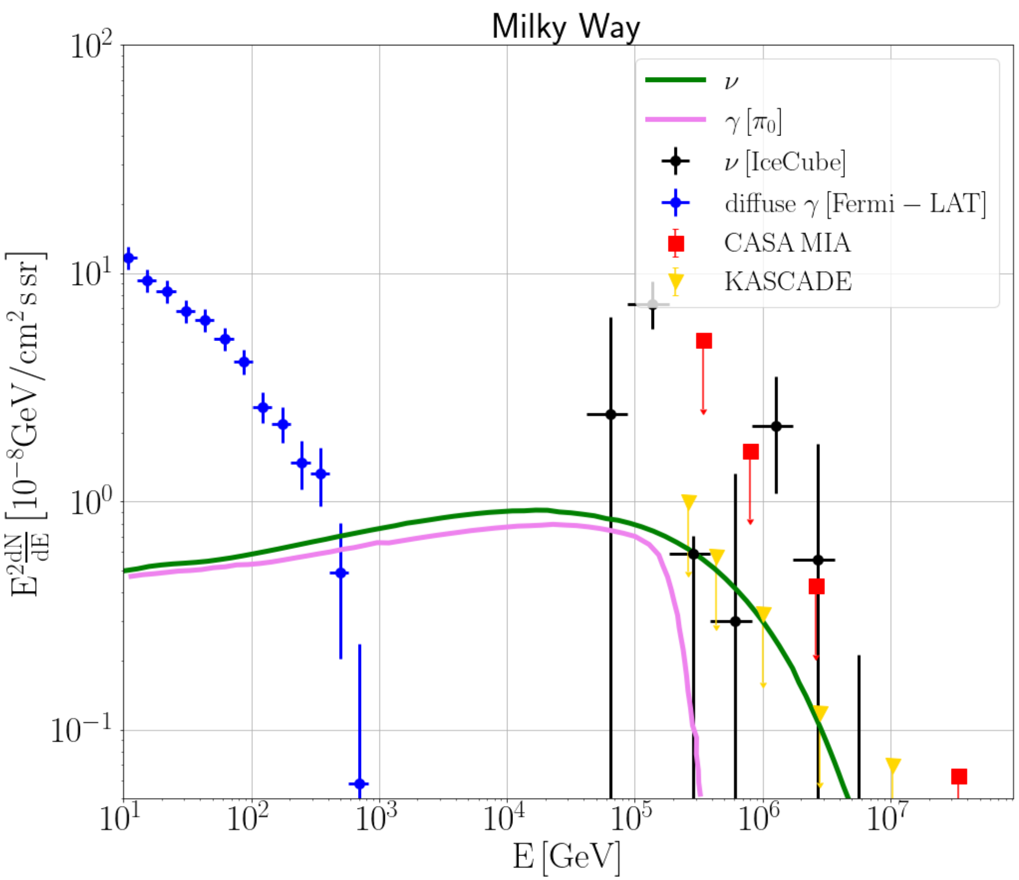
halo~10 kpc
problems with gamma?

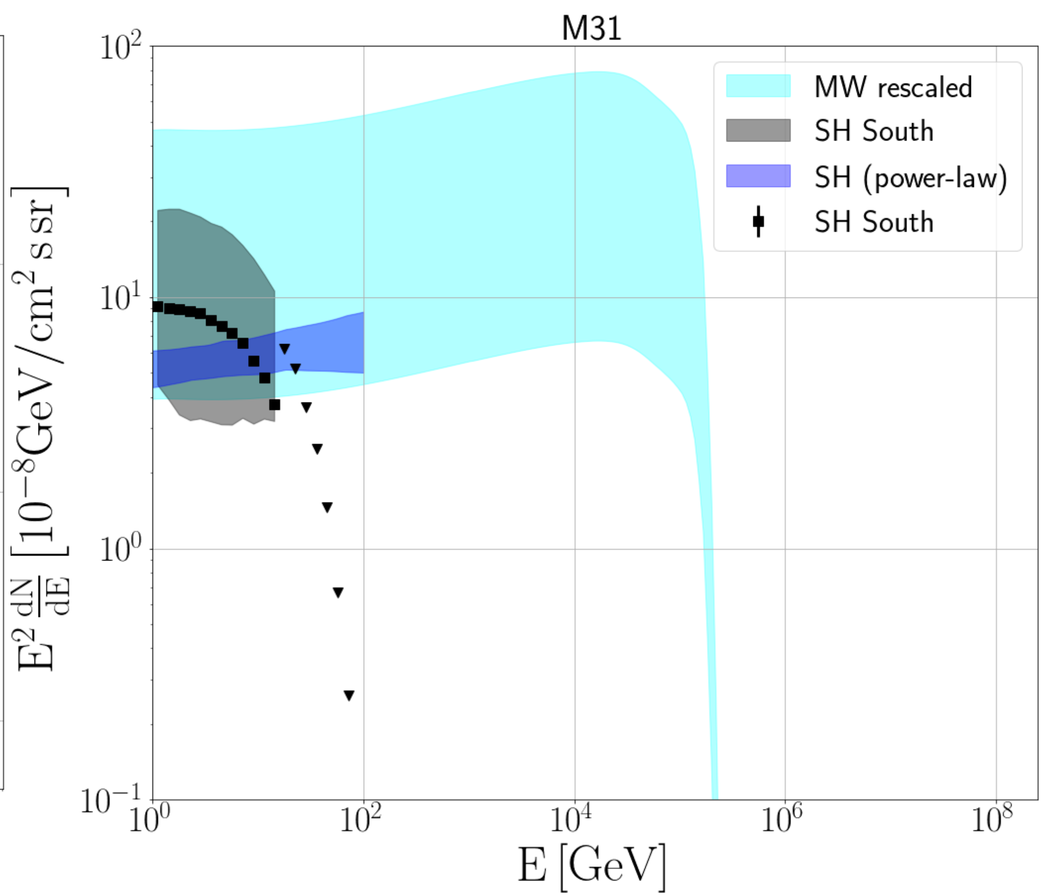
M31 - MW
gamma rays - neutrinos

Recchia et al. 2019

Assume diffuse Icecube neutrino flux
are produce in the MW halo
R~100 kpc

Rescale CR density
assuming that the emissions
in M31 and MW have a similar origin





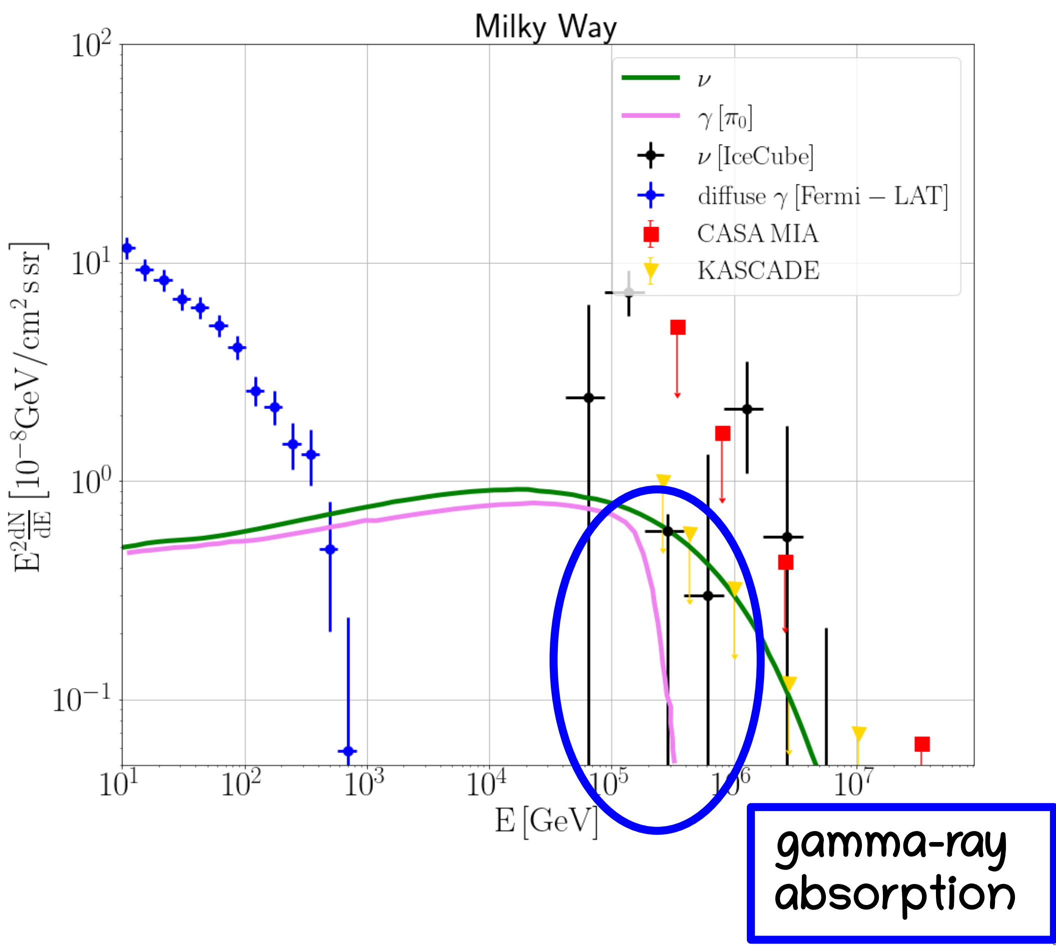
M31 - MW

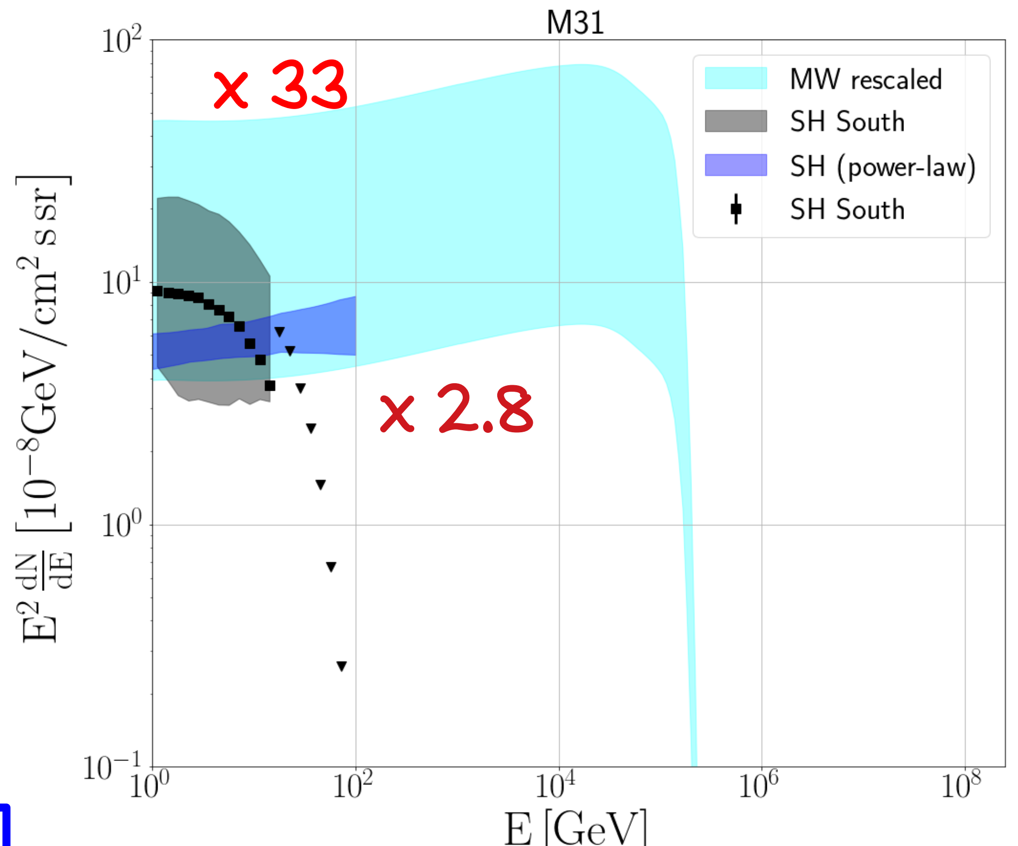
gamma rays - neutrinos

can work if hard gamma-ray spectrum in SH of M31

$(M_{M31}/M_{MW})^{3/2}$ → factor 2.8

$M_{BH,M31}/M_{BH,MW}$ → factor 33

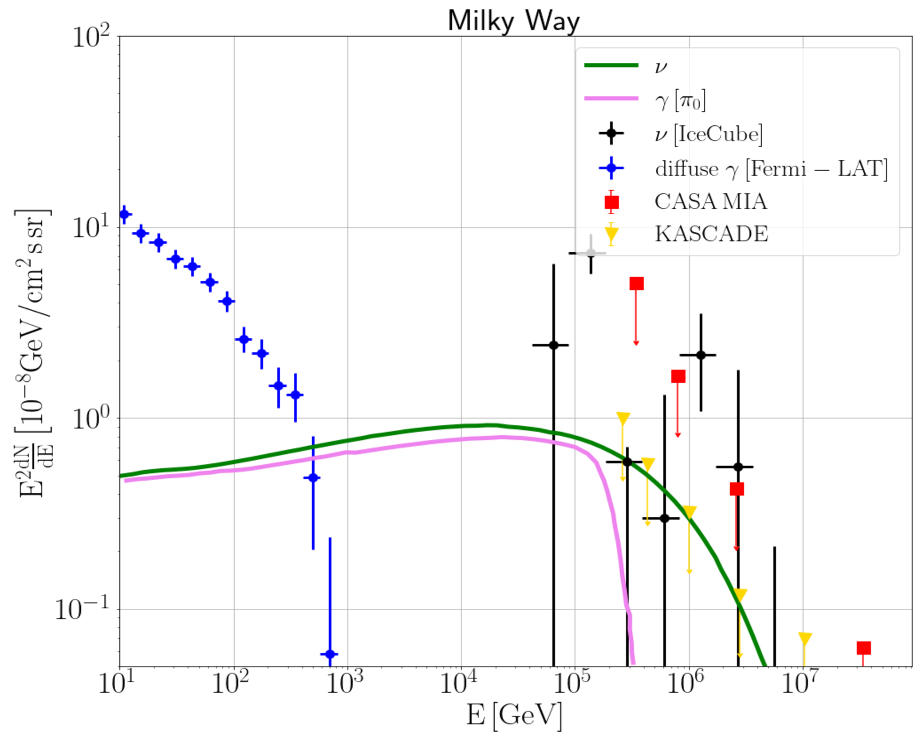




Vernetto & Lipari (2016)

M31 - MW: multimessanger

- recovery of **isotropic diffuse gamma**-ray emission **above ~TeV**
  - ~ 10^{-8} GeV/cm^2 s sr
  - up to < PeV



- M31 contribution to Icecube isotropic flux
  - about ~ 5%
  - extended ~ 15°

- contribution from MW-like galaxies within ~ 3 Gpc subdominant
  - flux from M31-like galaxy at d> 10 Mpc below detection limit
  - but **enhanced nuclear/starburst** (NGC 1068)

- **diffuse X-ray** from secondary e- << extragal. diffuse X-ray bck
  - **L^{MW}_x (1-100 keV) ~ 10^{37} erg/s** $R_{H,2}{}^2$
  - flux ~ 6×10^{-14} $R_{H,2}{}^2$ $(d/3\mathrm{Mpc})^{-2}$ erg/cm^2 s
  - θ ~ 2° $R_{H,2}(d/3\mathrm{Mpc})^{-1}$
  - **difficult to detect?**

$$E_X^{\mathrm{synch}} \approx 2B_{-7}E^2_{e,\mathrm{PeV}}\mathrm{keV}$$

$$E_X^{\mathrm{CMB}} \approx 4E^2_{e,\mathrm{GeV}}\mathrm{keV}$$

Take-Home Message

The extended gamma-ray emission of M31 can accounted for by

GC activity
(buoyant bubble, termination shock) → Scale with BH mass

Accretion shock → Scale with galaxy mass

- if spectrum is hard and up to ~PeV
- if similar scenarios hold in the MW

the isotropic Icecube neutrino flux may be naturally explained with a production in the MW halo

Perspectives & Issues

- **showed the feasibility of shock and buoyant bubbles scenarios but details should be addressed for both MW and M31**
  - existence of accretion shock?
  - formation of termination shock?
  - strong shocks? acceleration?

  - buoyant bubbles in MW-like galaxies?
  - larger structure beyond Fermi Bubbles?

- **is the gamma-ray spectrum of M31 hard up to ~PeV or cut-off at ~10-100 GeV?**
  - very important for interpreting Icecube neutrino
  - may be checked by LHAASO?

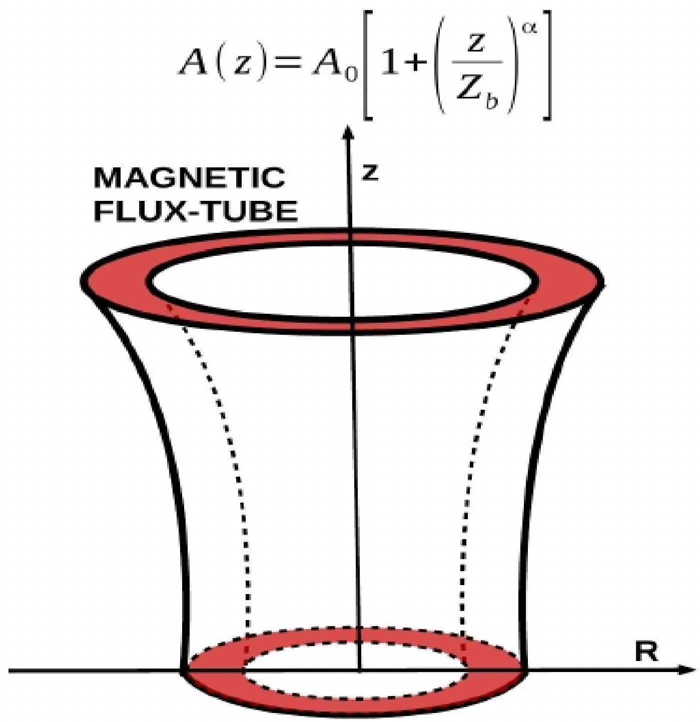
- **source of CR scattering at ~ 100 kpc in galactic halos?**

- check potential gamma-ray signals from LMC
- UHECRs from M31? Zirakashvili et al. 2022


Thank You

Propagation in galactic wind

Recchia et al. 2016

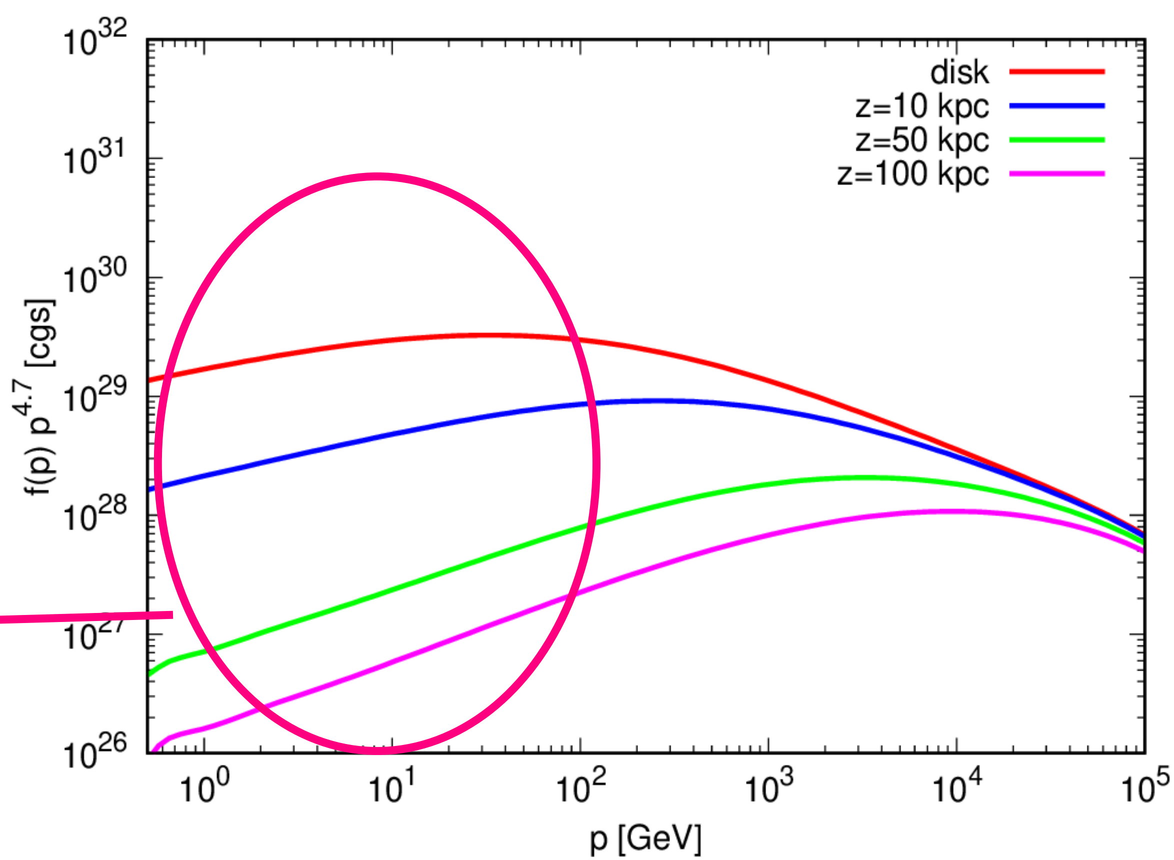


GALACTIC DISK

$$\epsilon(z) \propto n_{CR}(z) n_H(z)$$

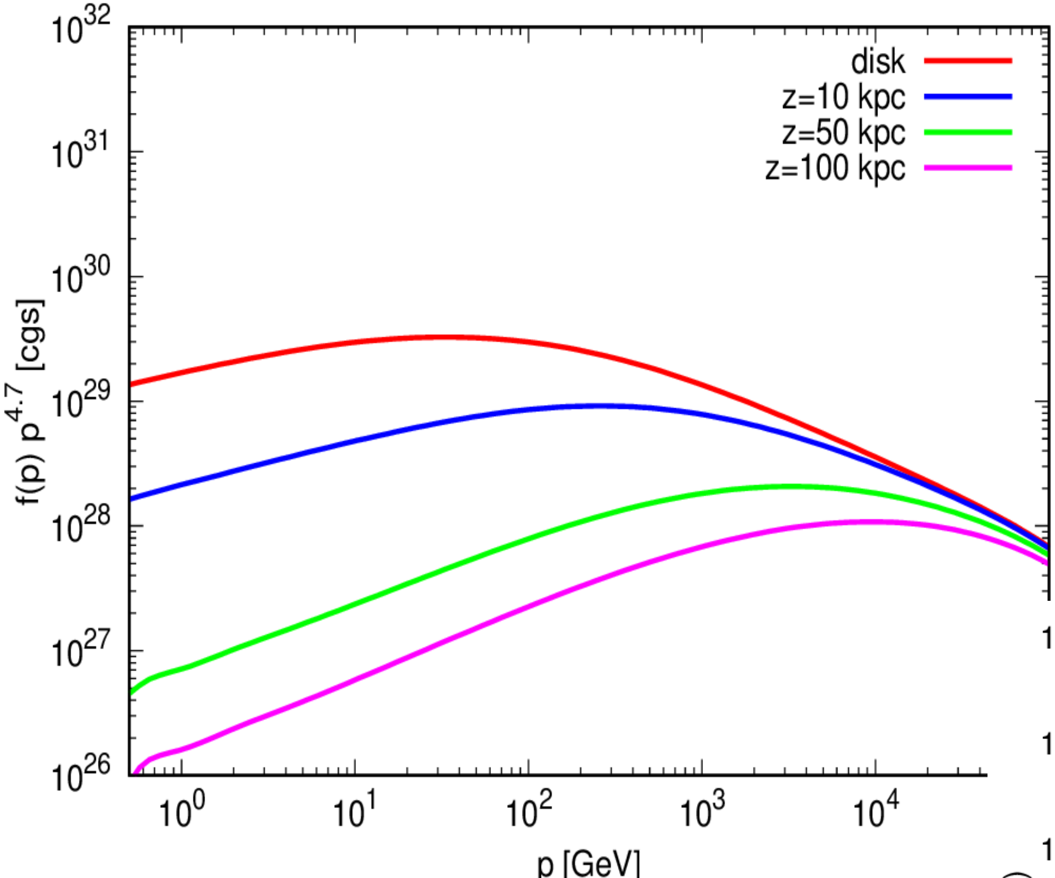
$D(z)$ diffusion coefficient

$u(z)$ galactic wind

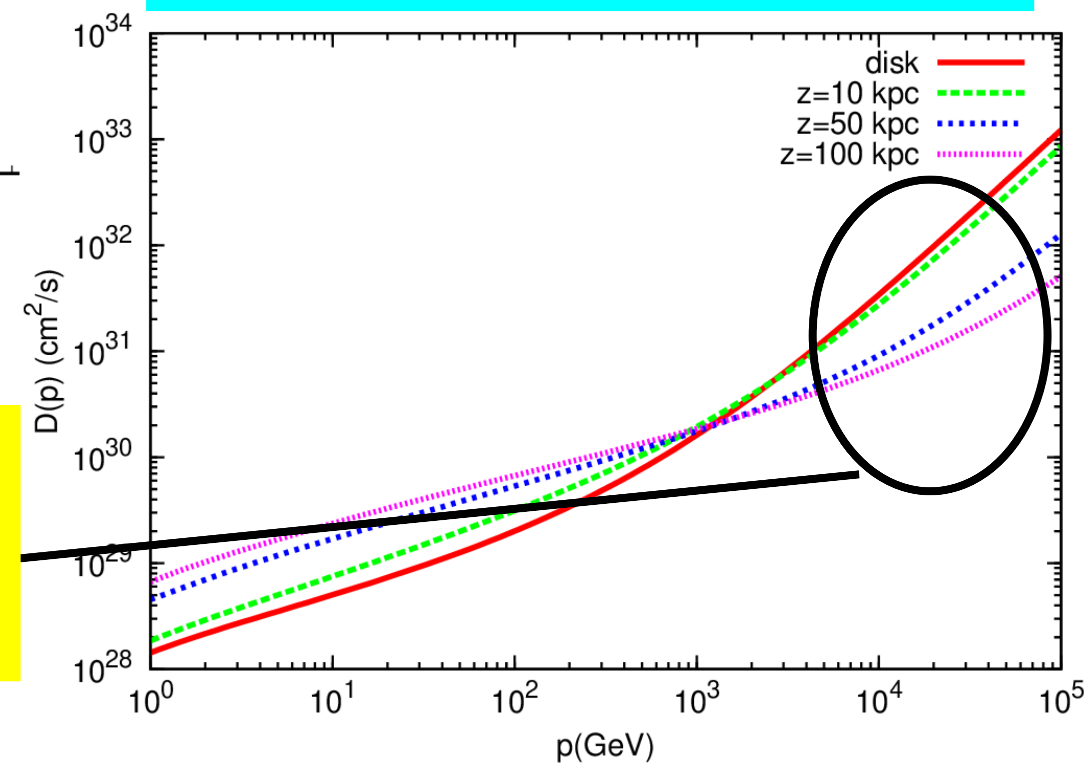


CR density
Strong decrease with z...

Propagation in galactic wind



Self-generated diffusion



effect of non linearity may give

- mild increase below ~ 1 TeV
- even a decrease above TeV